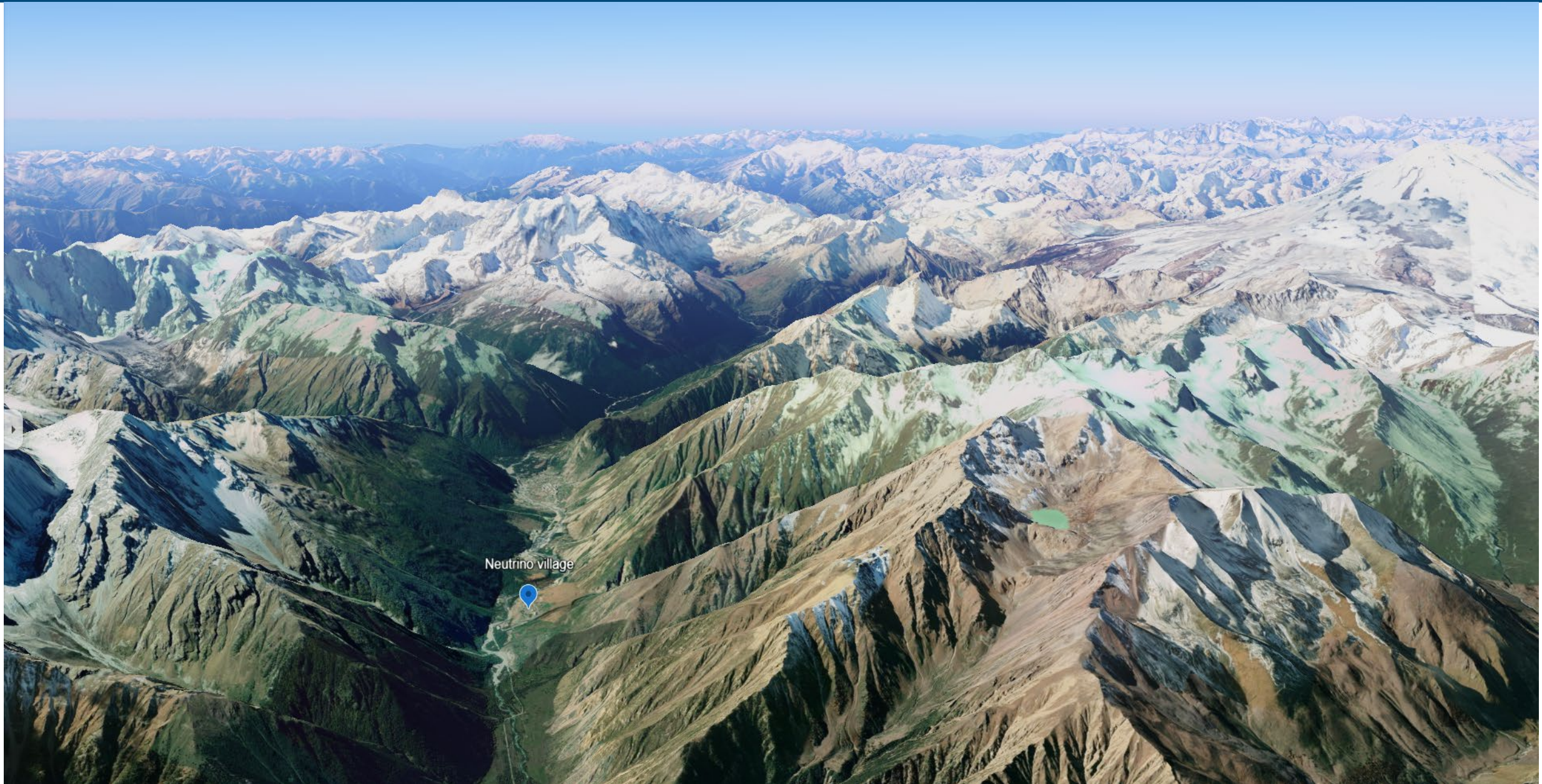




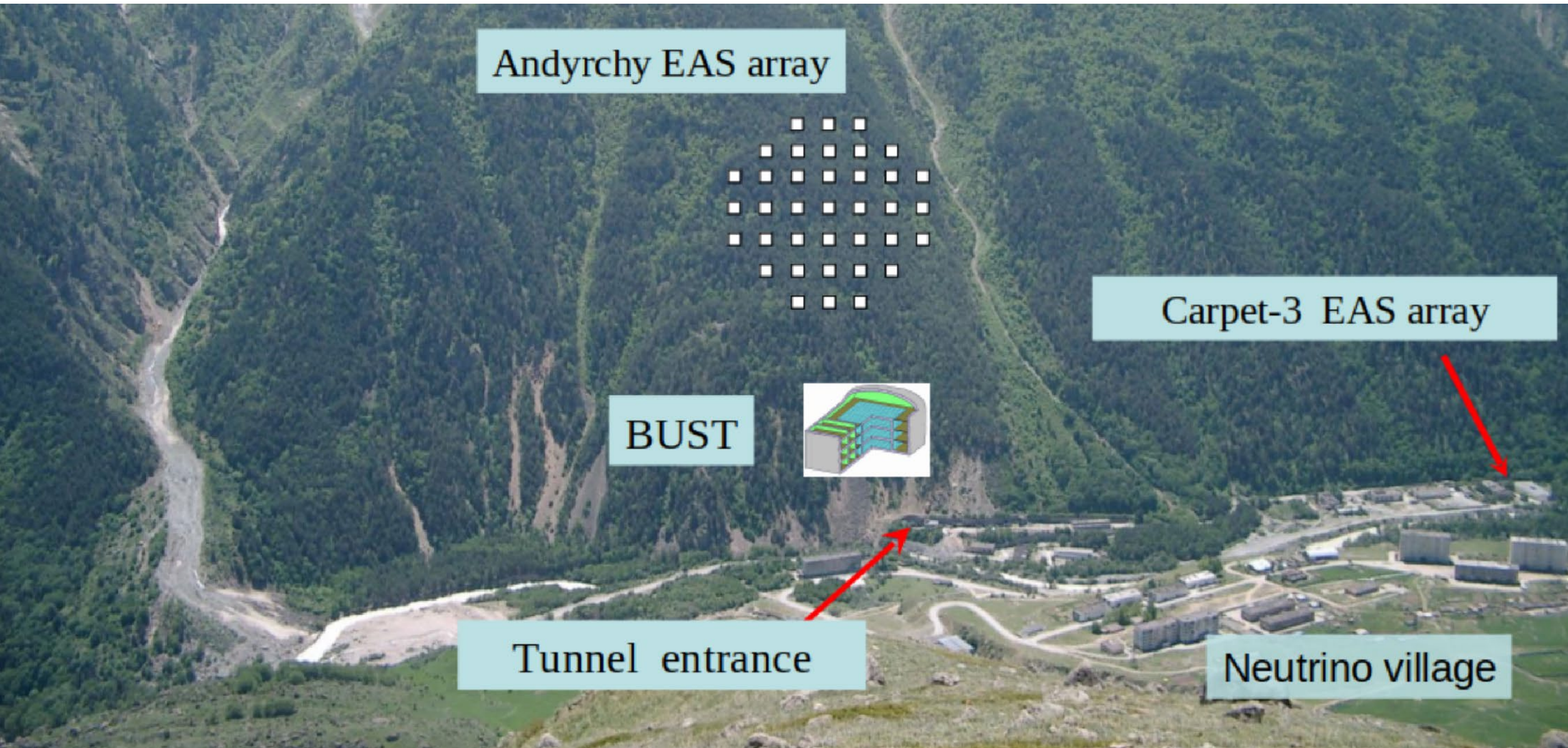
Recent results of the Carpet-3 collaboration

Nikita Vasiliev
from the name of the «Carpet-3» collaboration

Baksan Neutrino Observatory, Neutrino village

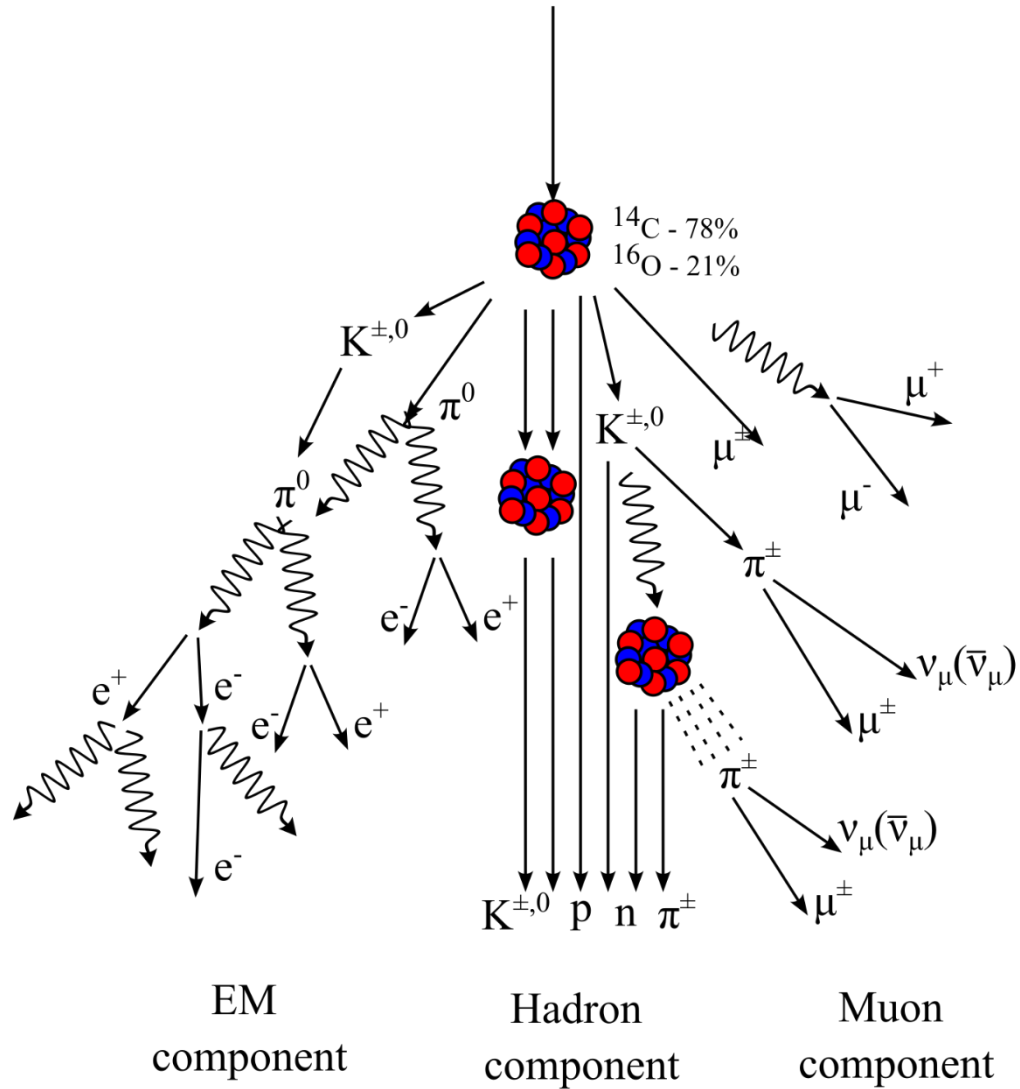


Baksan Neutrino Observatory detector complexes

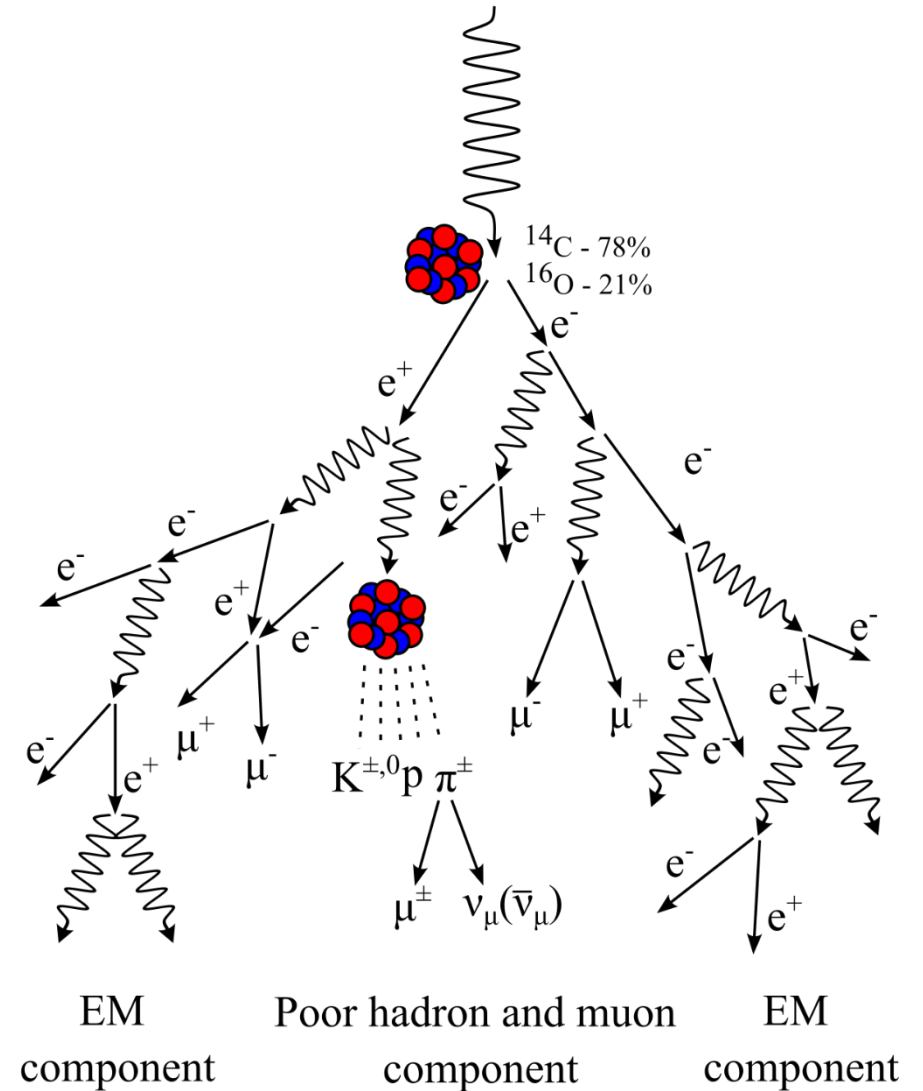


Extensive Air Showers (EAS)

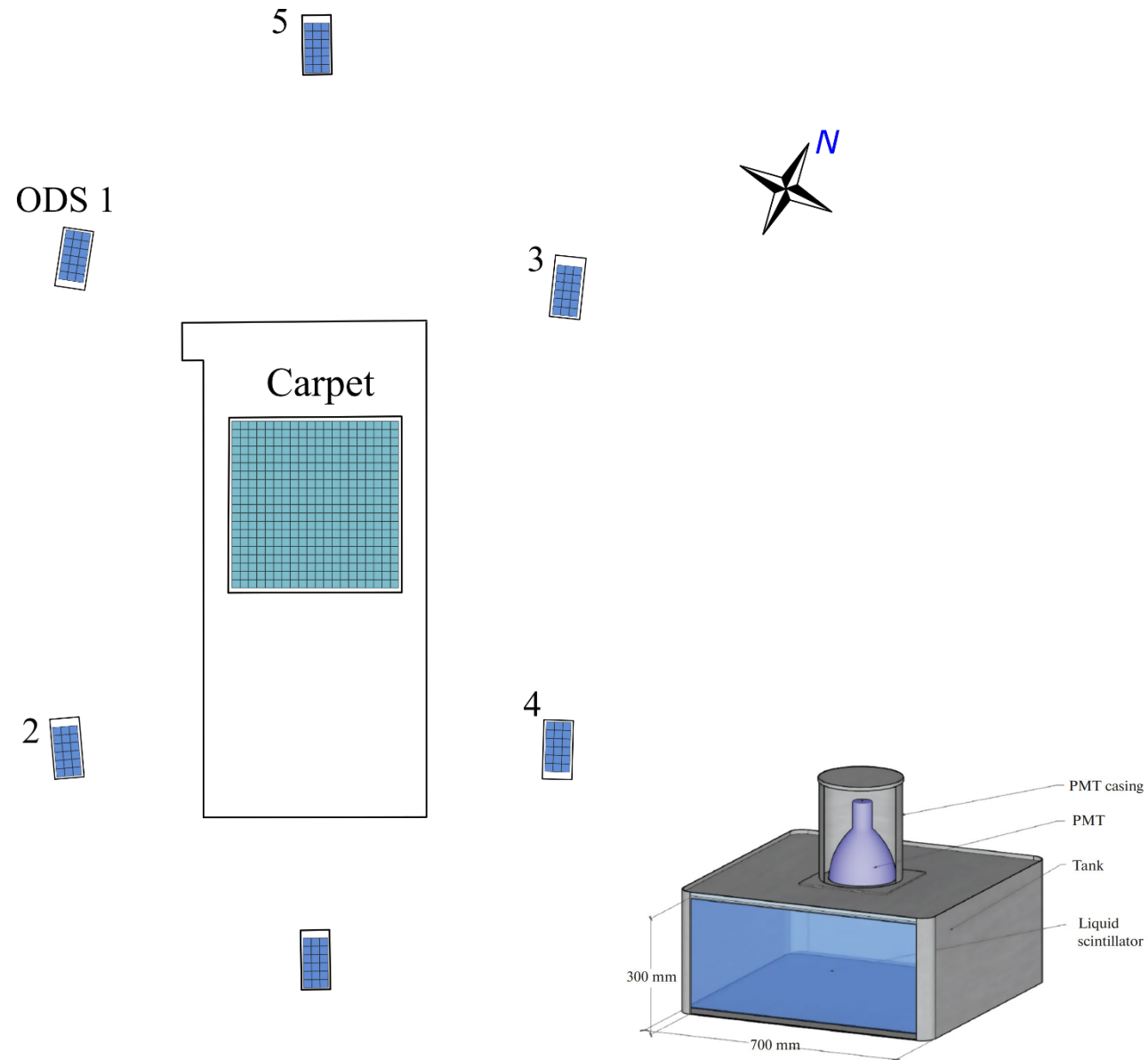
Primary particle: hadron
(p, α , ..., Fe)



Primary particle: photon

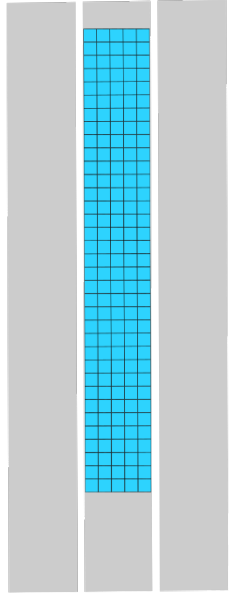


Carpet EAS array since ~1973



First modernization (Addition of the Muon Detector 175 m²)

Underground
Muon Detector

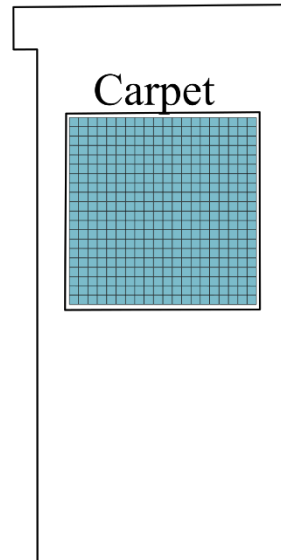


194650
482850


ODS 1


194650
482850

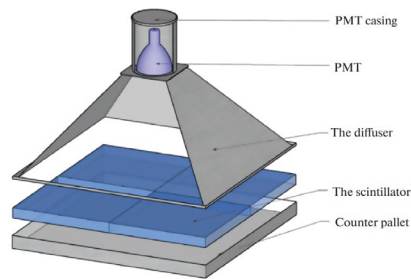
2

3


194650
482850
4


5

Carpet-3 facility (Current state)

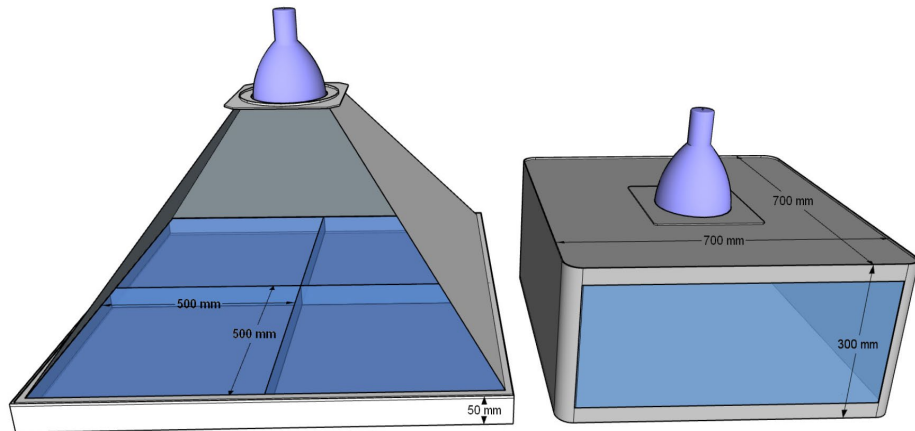
A – Carpet detector array;

B – outer detector stations (ODS);

C – Underground muon detector (MD);

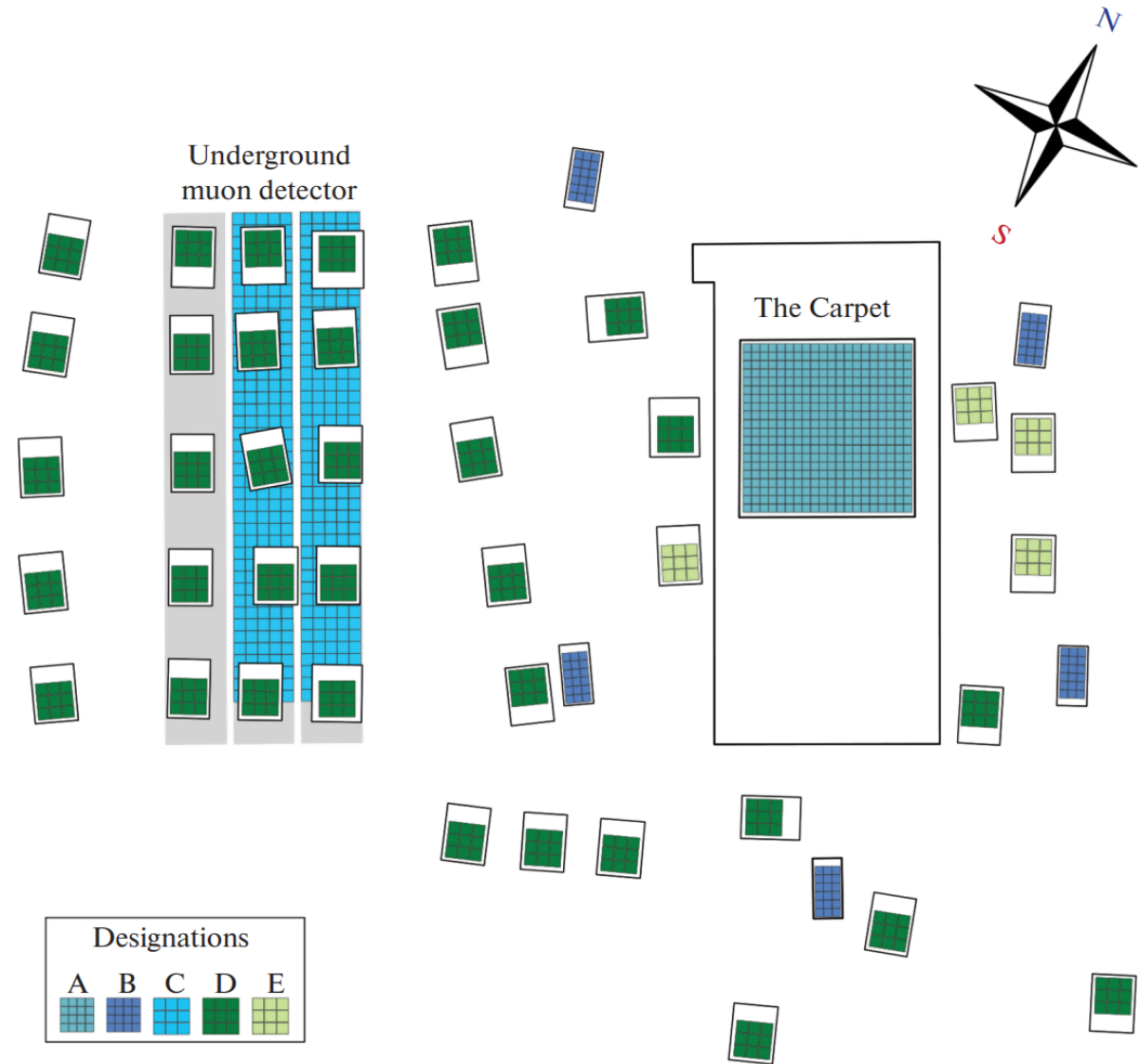
D – new ODS (plastic scintillator);

E – unfinished ODS;



1 r.p. ~ 10 MeV

1 r.p. ~ 50 MeV



r.p. (relativistic particle) – most probable energy deposition in the scintillator created by relativistic muons.

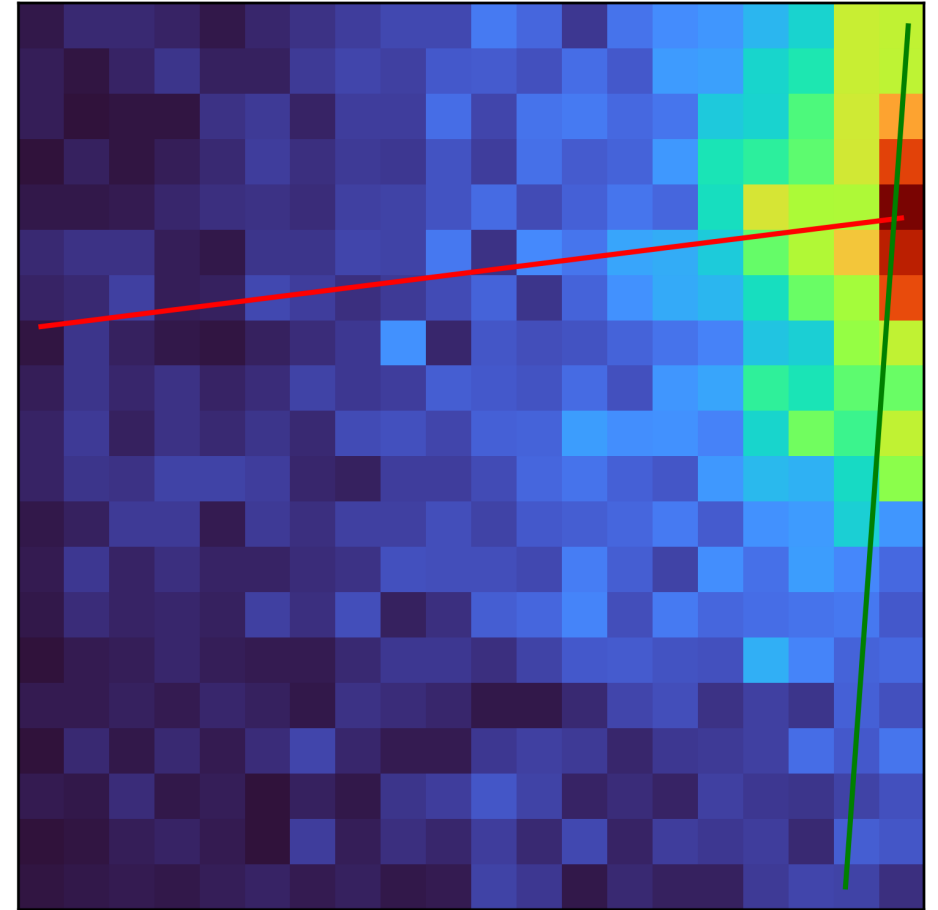
Reconstruction methodology

- ❖ EAS arrival direction (θ , φ) is reconstructed by relative time lags of 4 main ODS with flat shower front approximation.
- ❖ Shower size N_e and age s is reconstructed by fitting particle density spatial distribution with NKG function:

$$\rho_e(r, s, N_e) = \left(\frac{N_e}{r_M^2} \right) \frac{\Gamma(4, 5 - s_N)}{2\pi\Gamma(s_N)\Gamma(4, 5 - 2s_N)} \left(\frac{r}{r_M} \right)^{s_N-2} \left(1 + \frac{r}{r_M} \right)^{s_N-4,5}$$

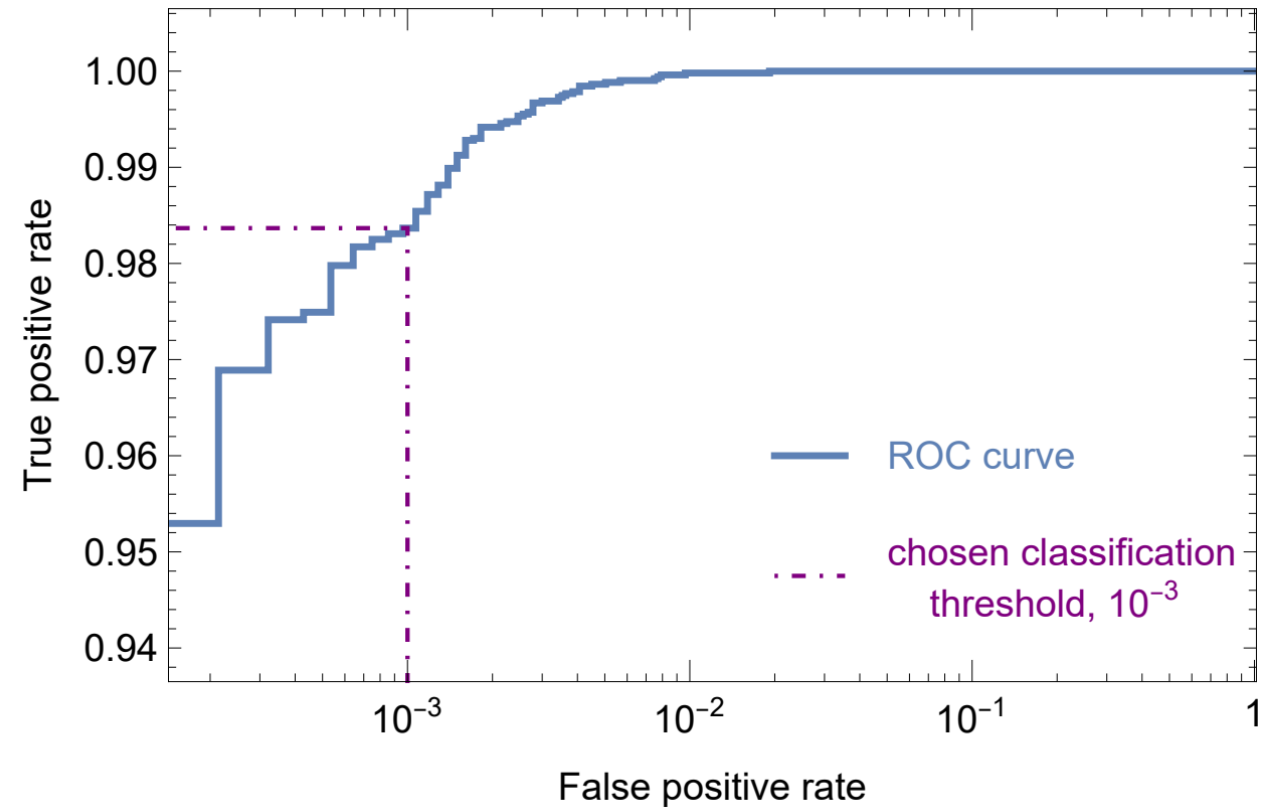
Shower core position reconstruction

- Identify the position of the detector with maximum energy deposition in each of the 20 central array rows/columns.
- Assign a weight equal to the total energy deposition in the row/column.
- Fit straight lines to the weighted points via LSM.
- The intersection of the lines defines the shower axis.

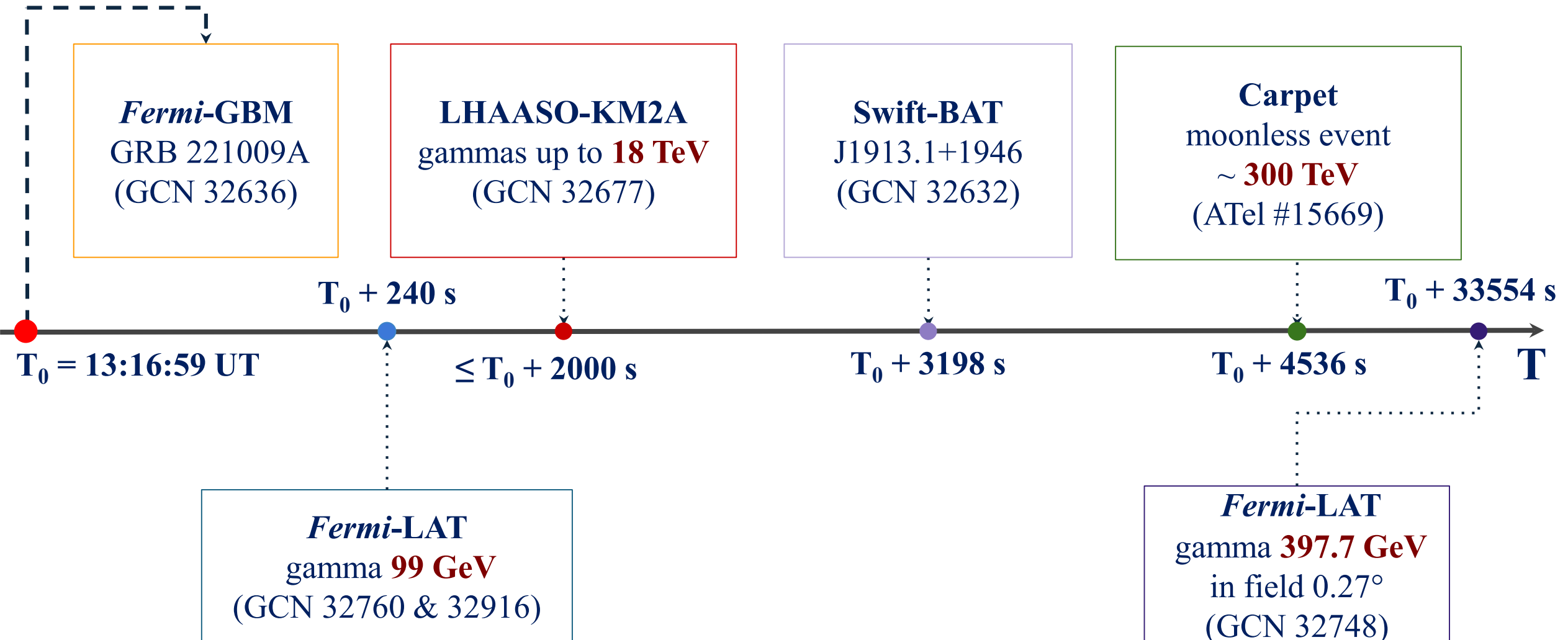


Neural network gamma-ray classification

- ❖ We estimate the type of primary particle using a neural network classifier trained on the MC event set (80609 events).
- ❖ The network is trained to distinguish between events with proton and photon primary particles.

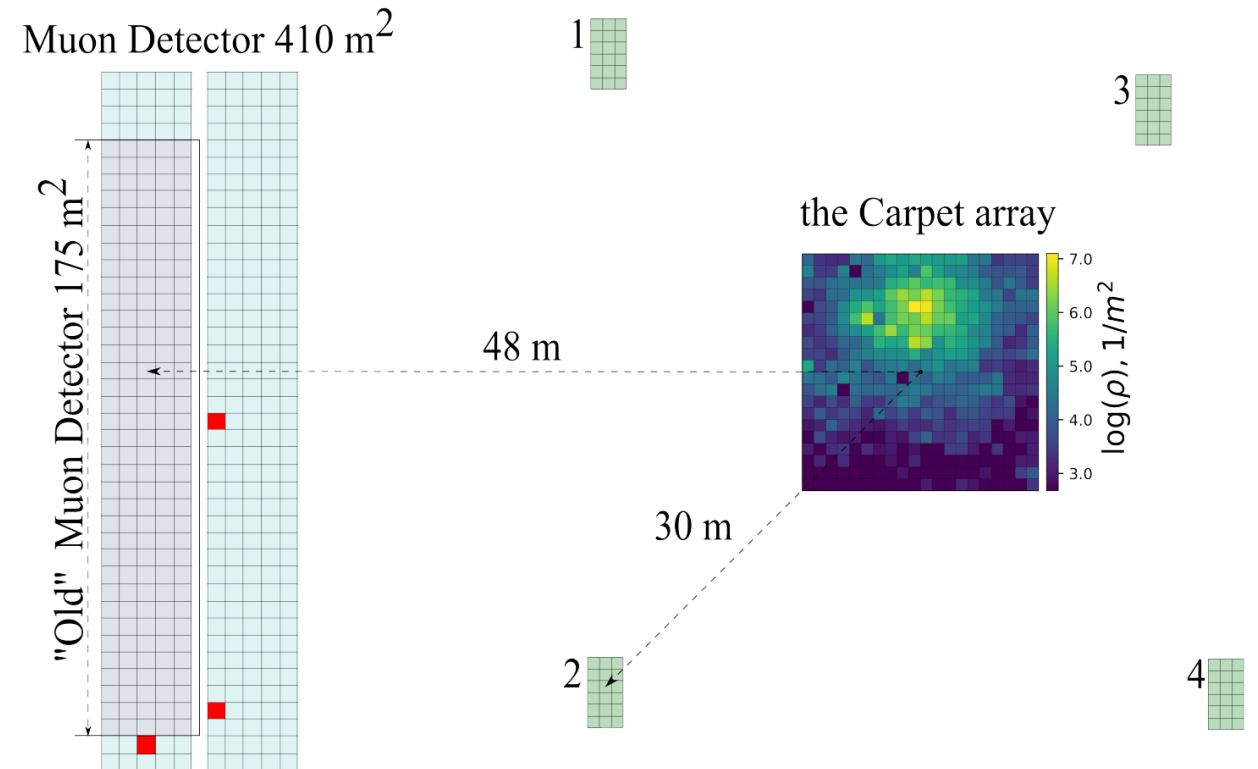


UHE γ -astronomy at the BNO (GRB221009A)

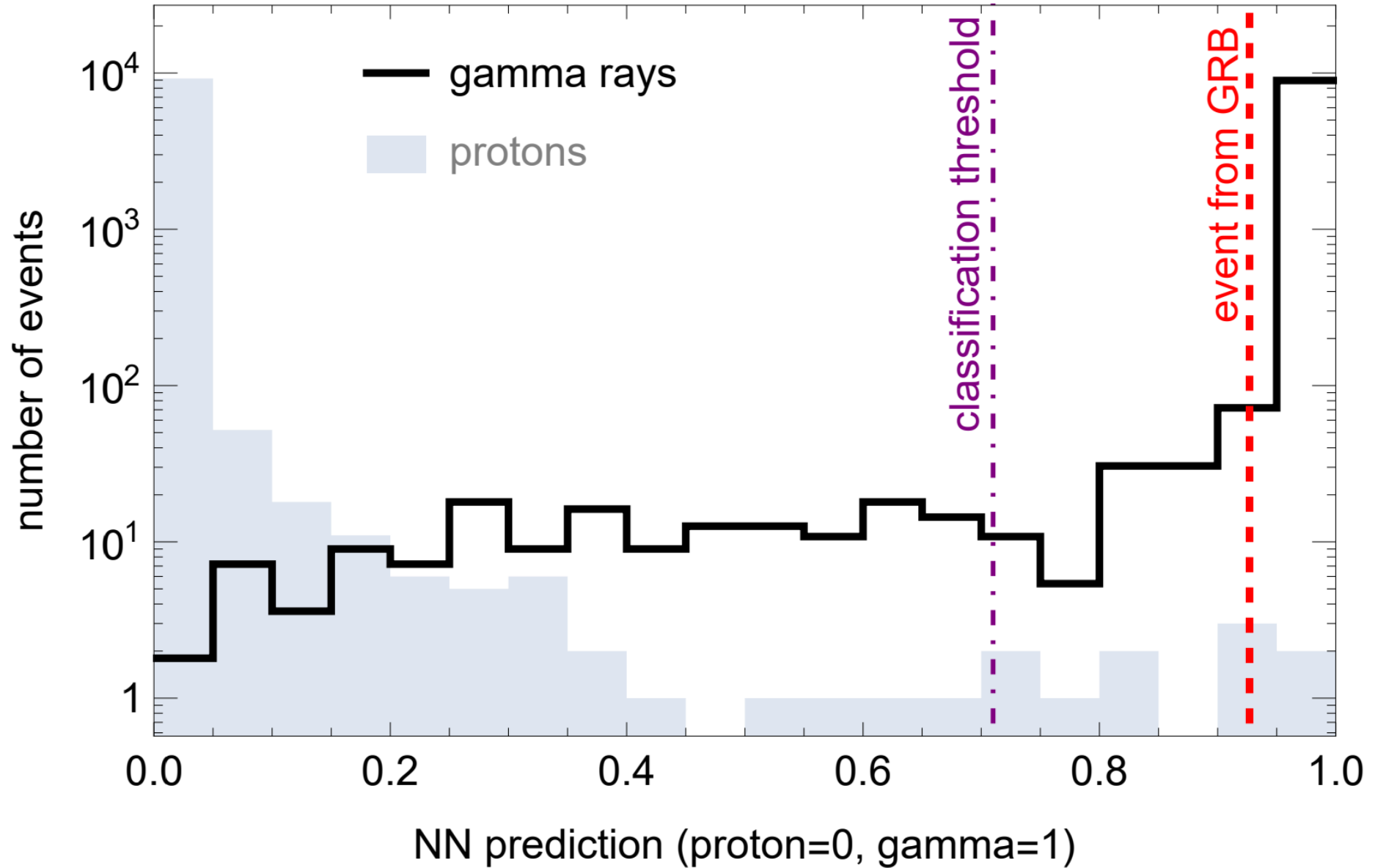


The event associated with GRB 221009A

- ❖ Estimated energy of the primary particle ($E=300 \pm 40$ TeV).
- ❖ Air Shower is photon-like (probability of a hadronic primary is 0,03 %).
- ❖ The event is coincident with the GRB in its arrival direction and time (chance probability $\sim 0,9$ %).

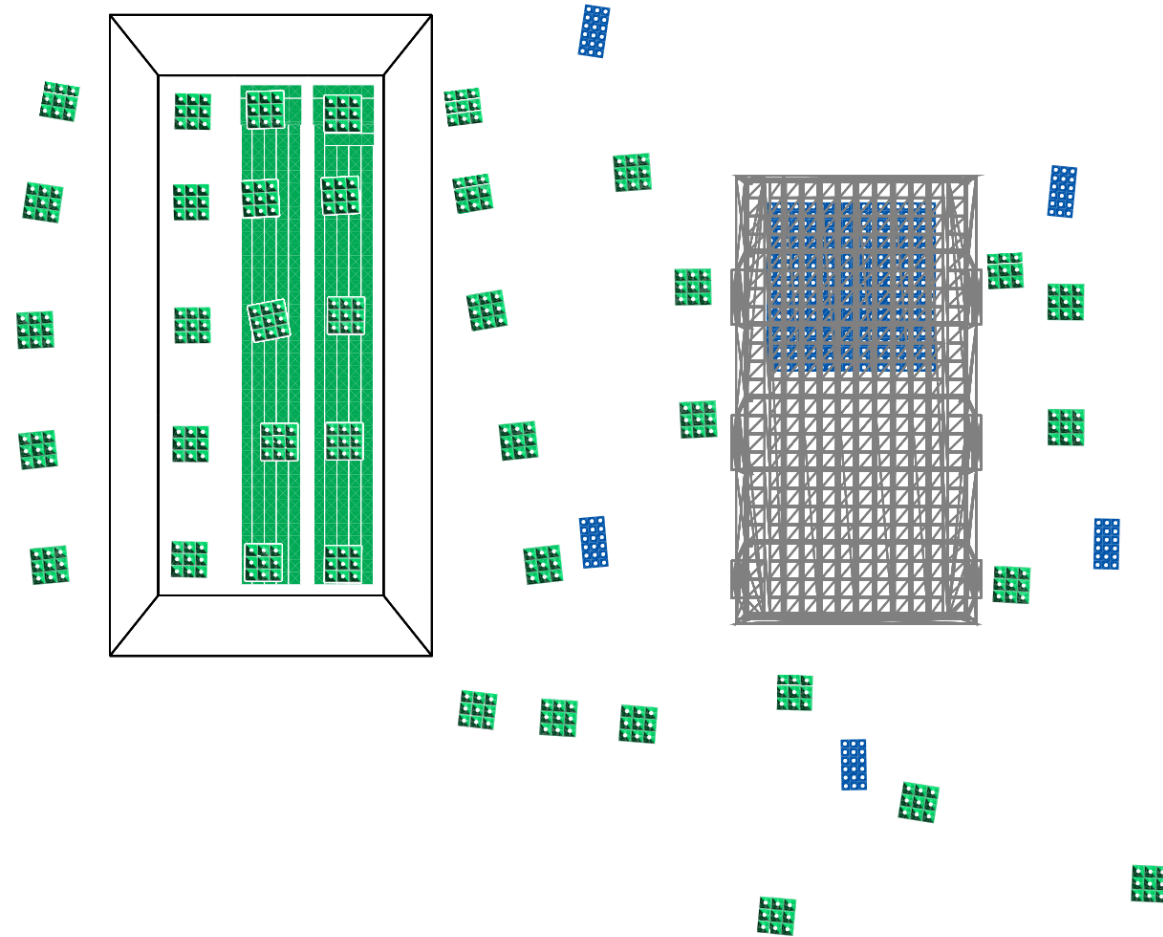


GRB event classification



Carpet-3 model in Geant4

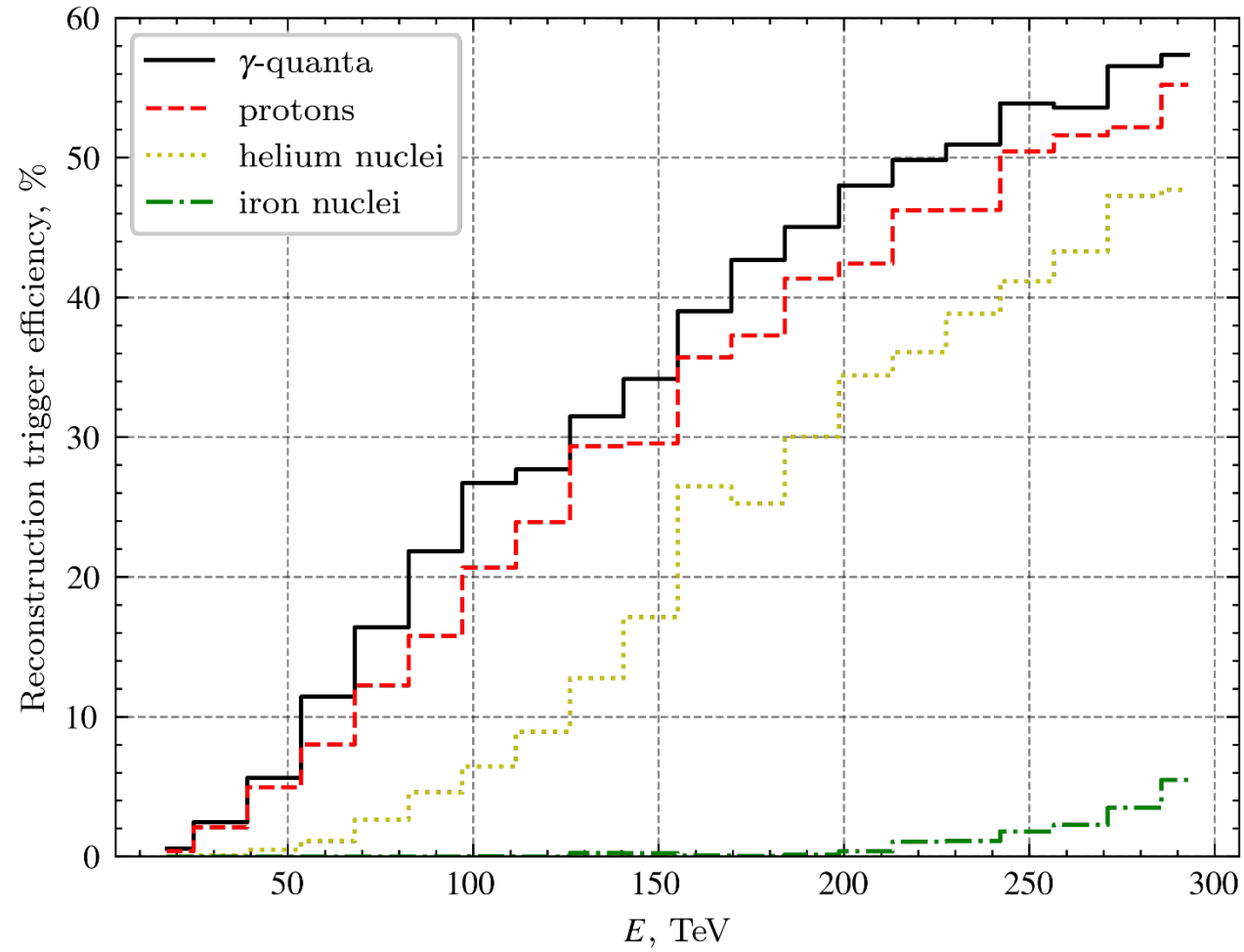
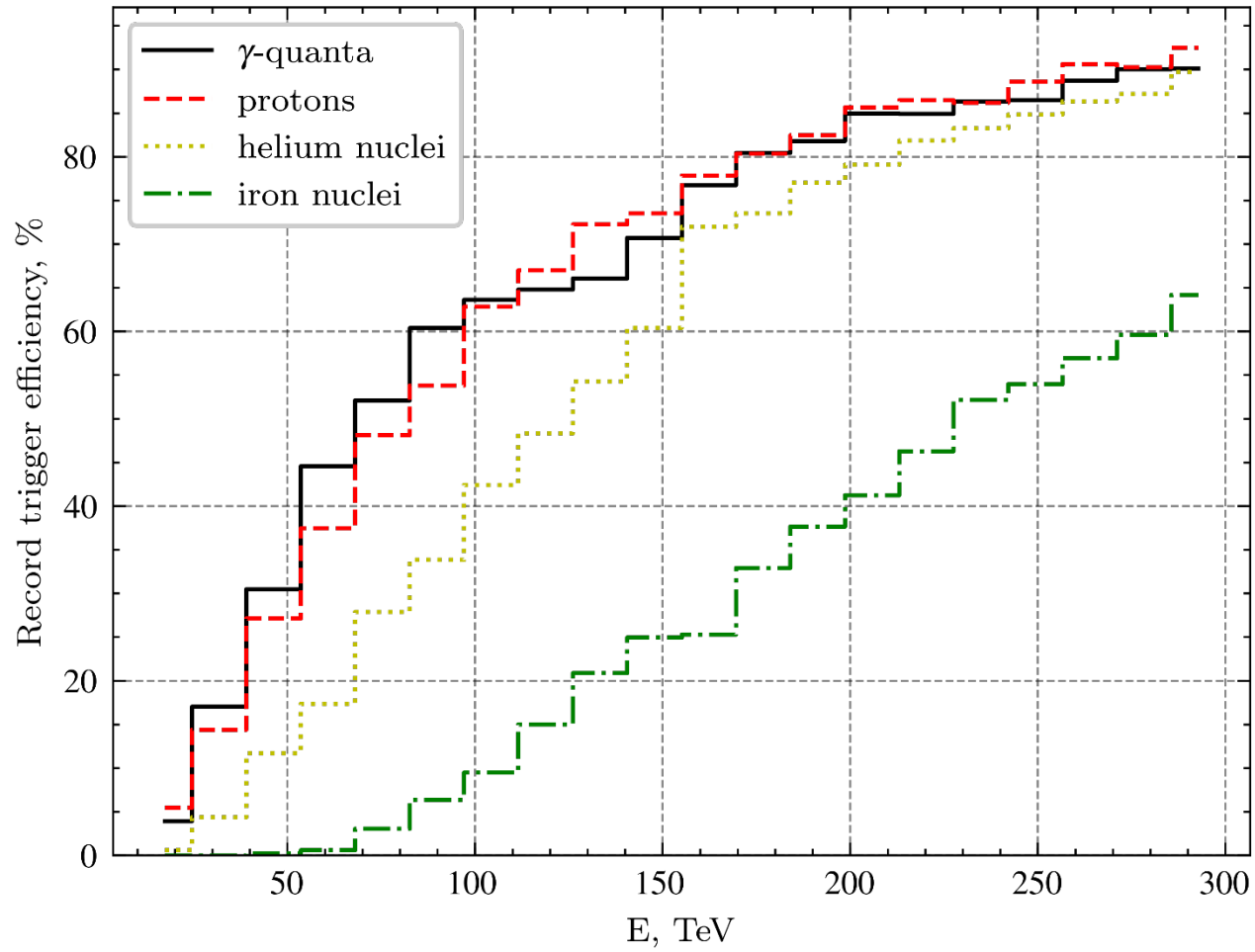
- ❖ Includes 1251 scintillation detectors.
- ❖ Passive geometry taken into account (central Carpet building and muon detector embankment).
- ❖ The response of each detector is calculated by simulating scintillation photons reaching the PMT photocathode.



Record and reconstruction triggers

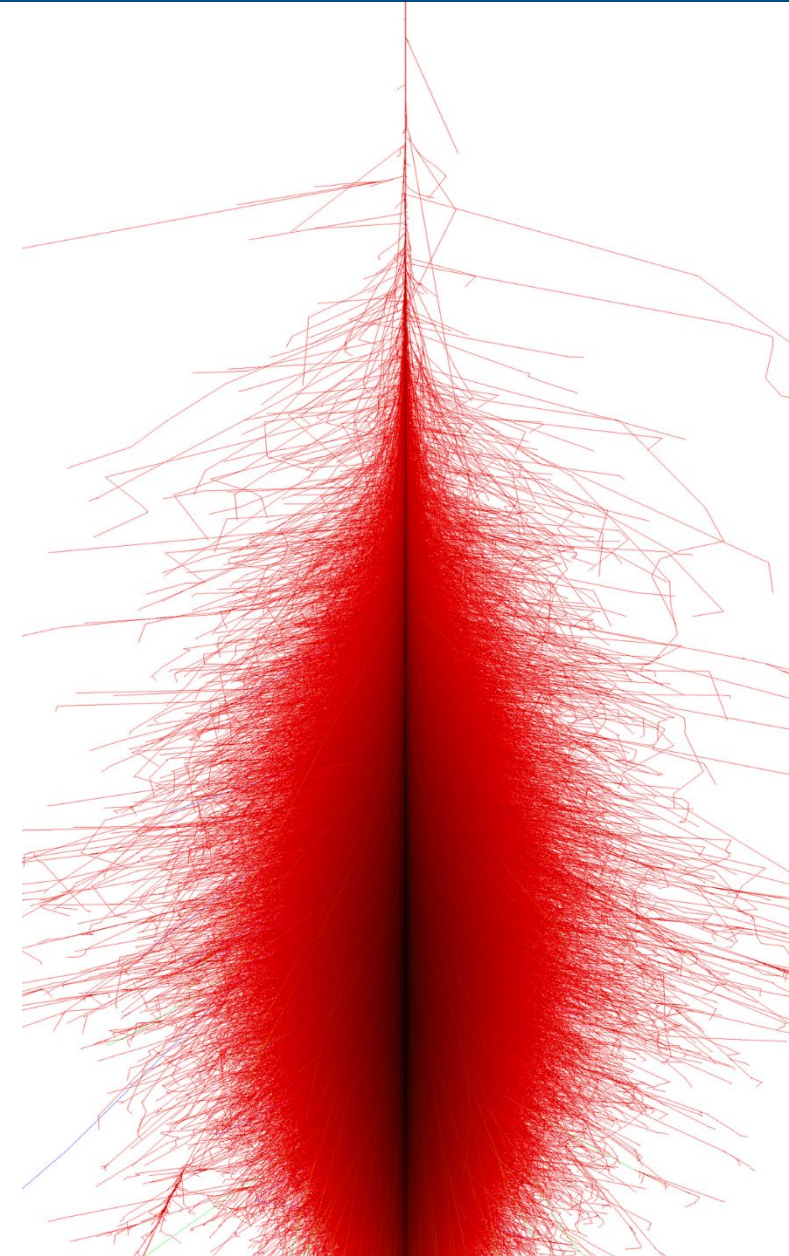
- ❖ Total energy deposition in the central Carpet exceeds 15 r.p.
- ❖ At least four ODS exceed the 0.5 r.p. threshold.
- ❖ Total energy deposition in the central Carpet exceeds 15 r.p.
- ❖ At least four ODS exceed the 0.5 r.p. threshold.
- ❖ At least 8 r.p. in 50 central cells
- ❖ Reconstructed axis must lie within the central $12.6 \times 12.6 \text{ m}^2$ area.
- ❖ Reconstructed zenith angle $\theta \leq 40$.

Trigger efficiency estimation



Conclusions

- ❖ A detailed description of the Carpet-3 detector array is introduced.
- ❖ A set of reconstruction algorithms of such parameters as EAS core position, arrival direction, age, size and primary type is presented.
- ❖ The developed methodology allowed us to estimate the efficiency of the detector array depending on the energy and the type of primary particle.
- ❖ Main goals for future research are developing more accurate reconstruction methods with the help of machine learning + testing the model precision on a large scale of energies.



Thank you for your attention!

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O. I. Mikhailova¹, V. B. Petkov^{1,2}, E. I. Podlesnyi^{4,1}, N. A. Pozdnukhov¹, V. S. Romanenko¹,
G. I. Rubtsov¹, S. V. Troitsky^{1,3}, I. B. Unatlov¹, N.A. Vasilev¹, A. F. Yanin¹, K. V. Zhuravleva¹**

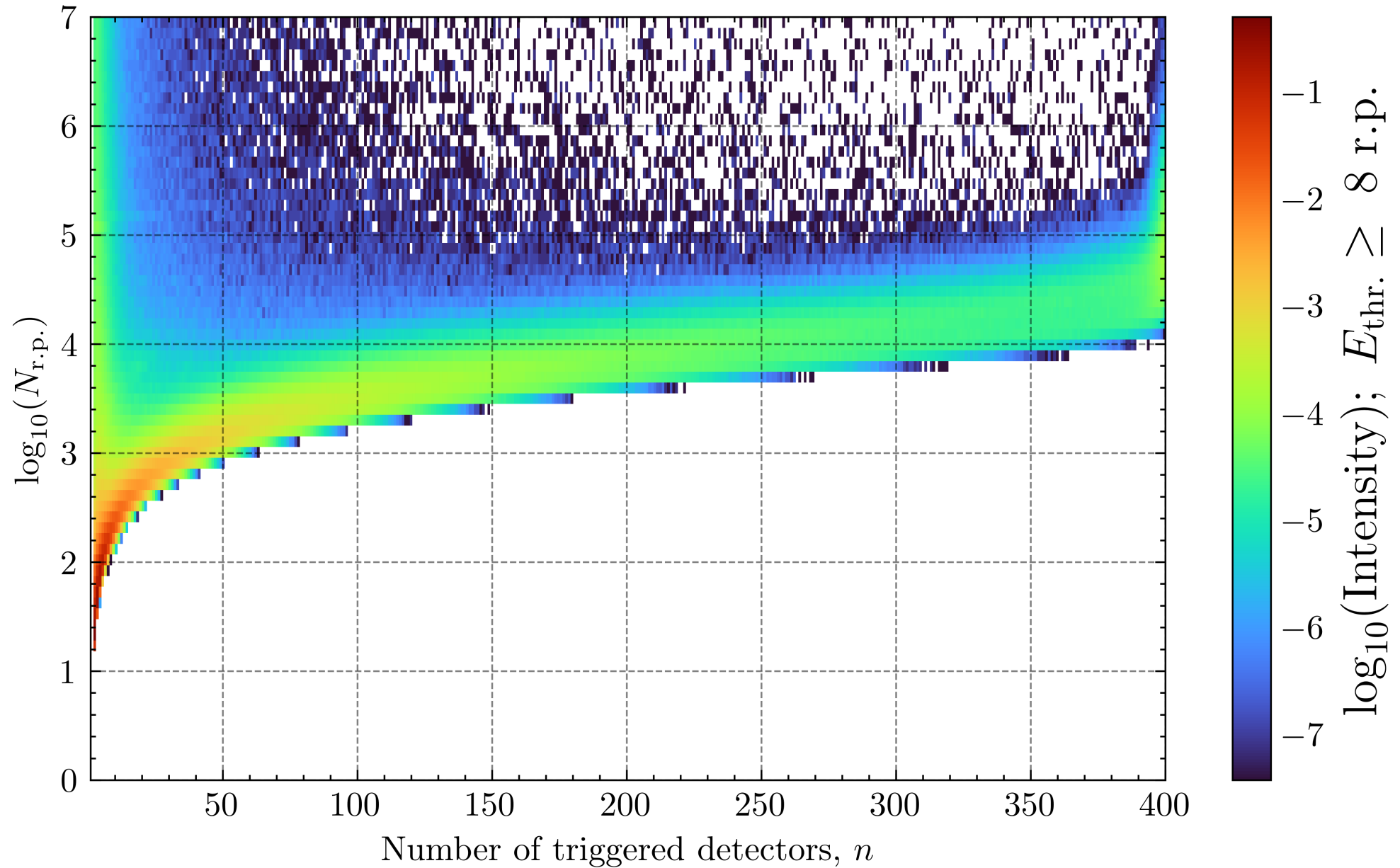
¹ Institute for Nuclear Research of the RAS, Moscow, Russia

² Institute of Astronomy, Russian Academy of Sciences, Moscow, Russia

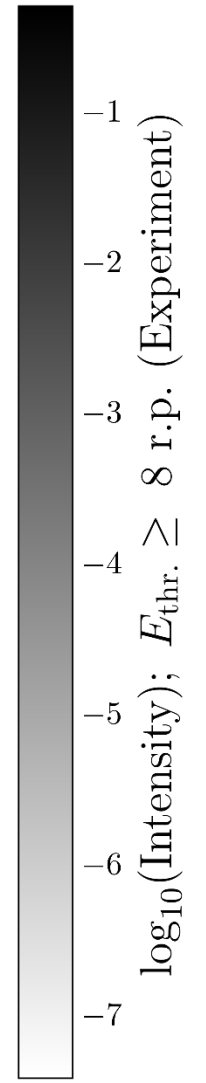
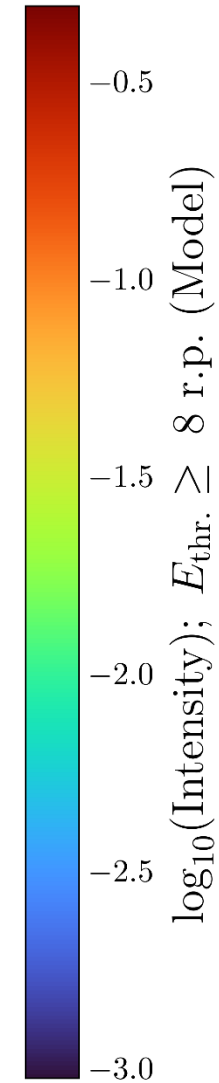
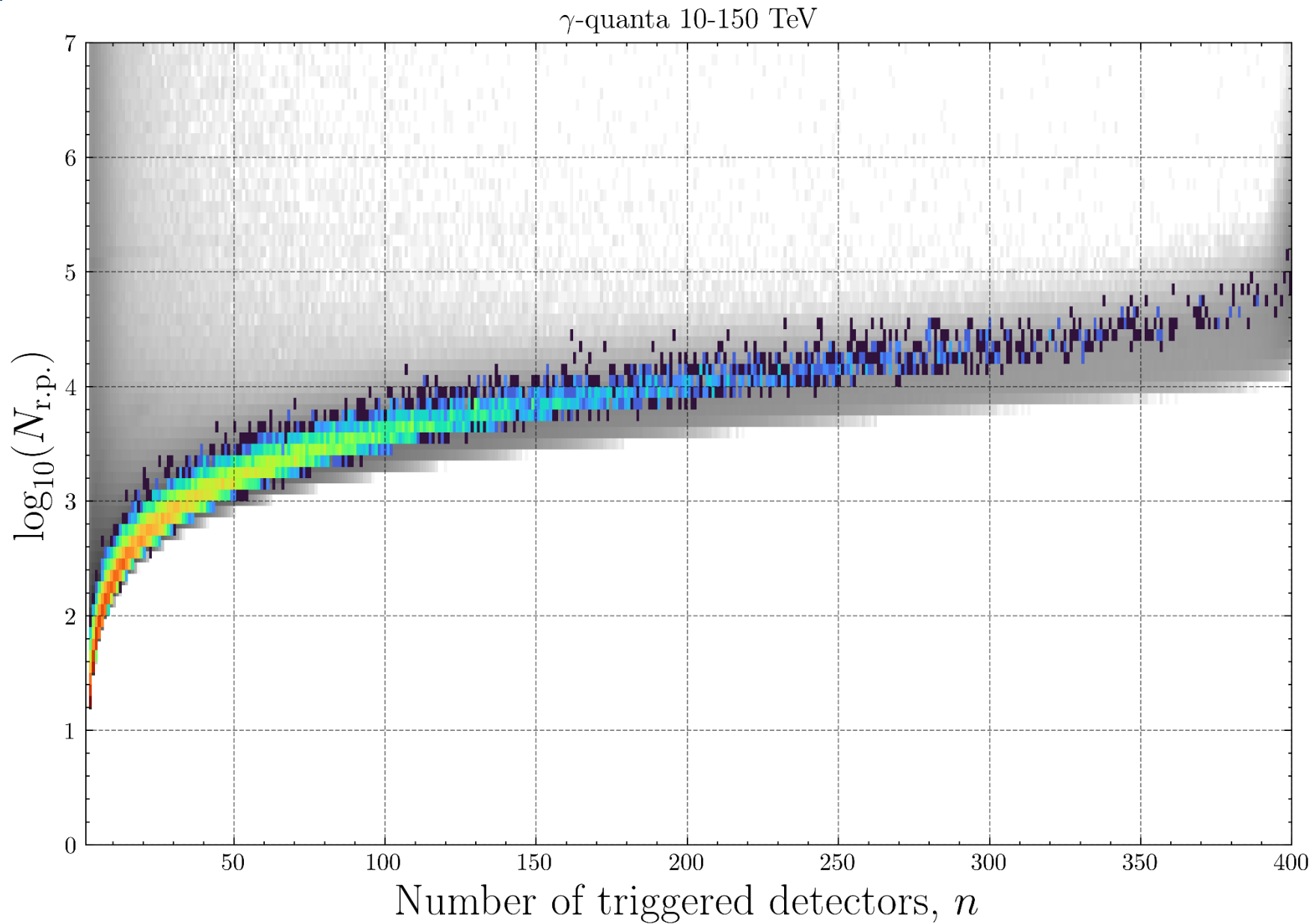
³ Lomonosov Moscow State University, Moscow, Russia

⁴ Norwegian University for Science and Technology, Institutt for fysikk, Trondheim, Norway

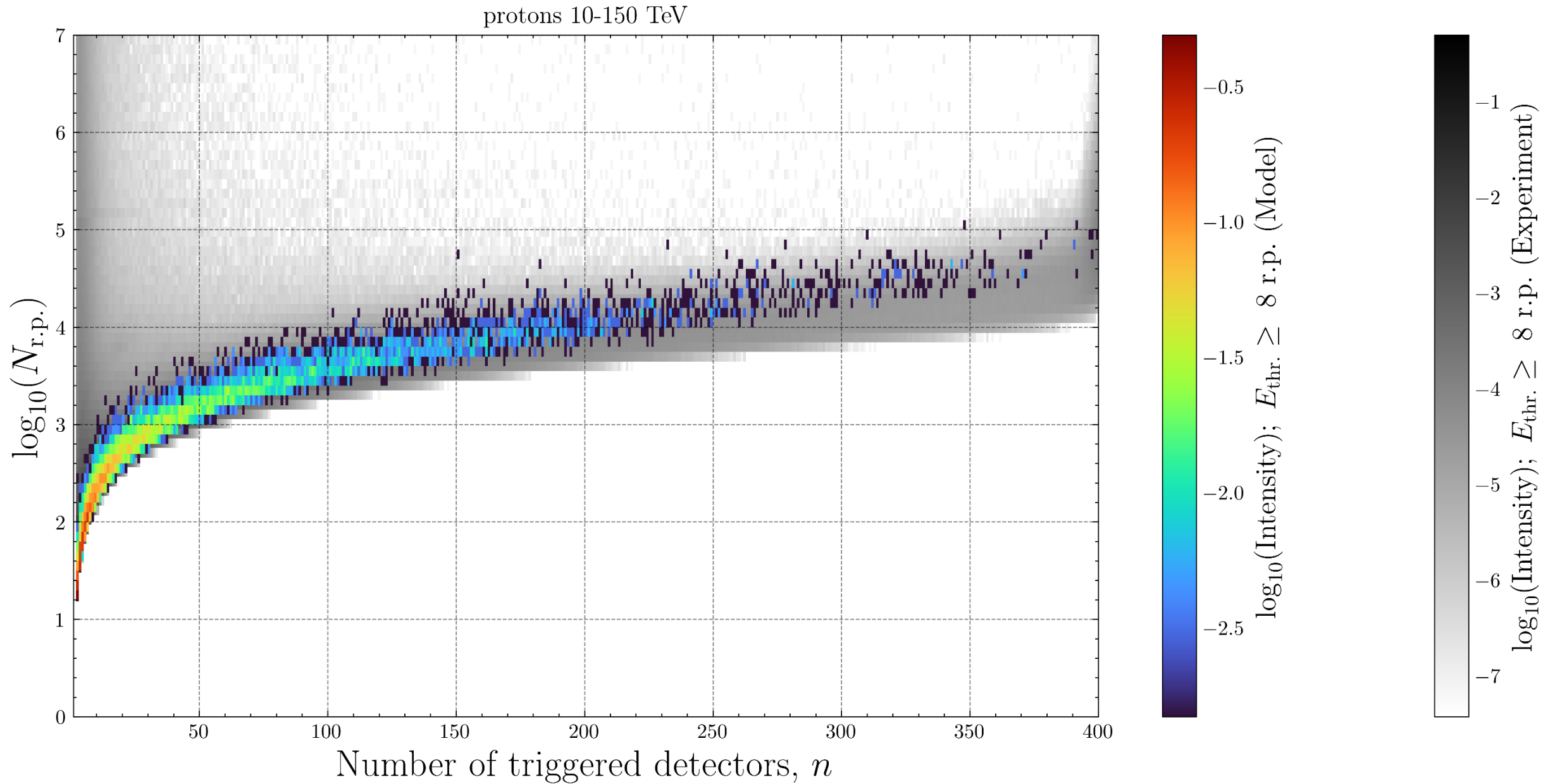
- ❖ Over 200000 EAS events were generated for primary γ -quanta, protons, helium, and iron nuclei.
- ❖ Zenith angles ranged from 0° to 40° .
- ❖ Discrete uniform energy distribution from 10 to 300 TeV (1 TeV steps).
- ❖ Observation level: 1700 m above sea level.



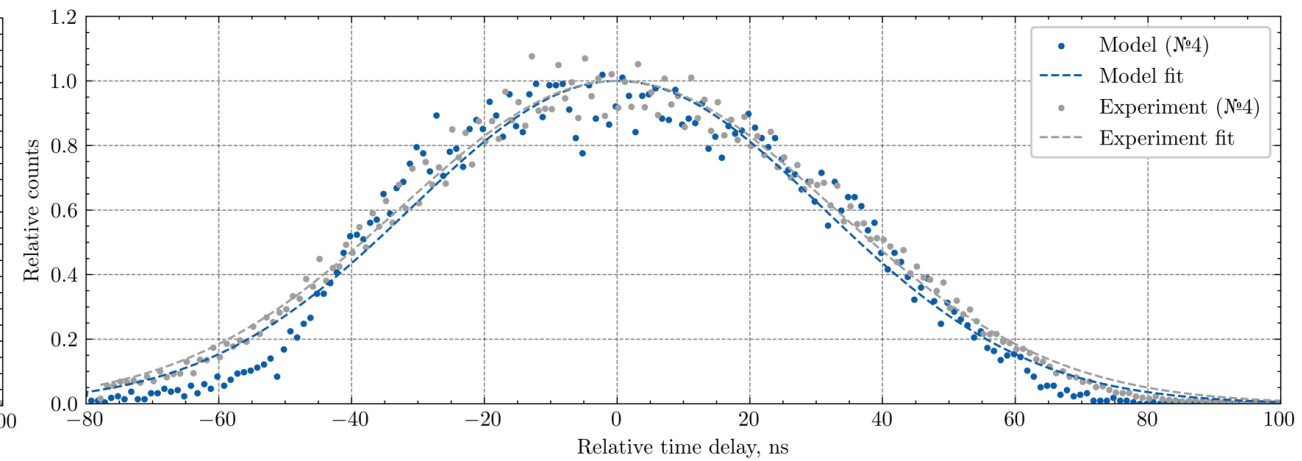
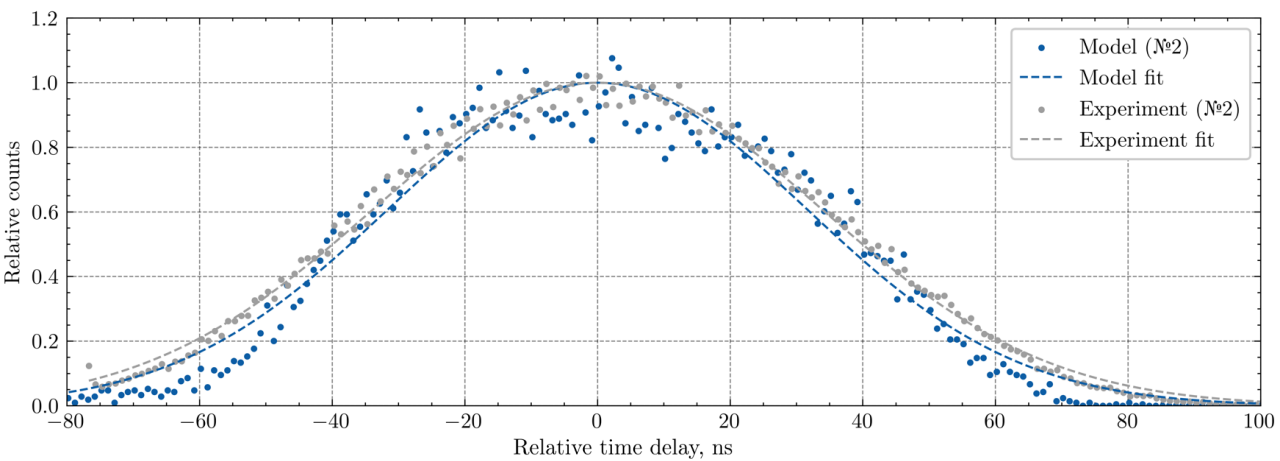
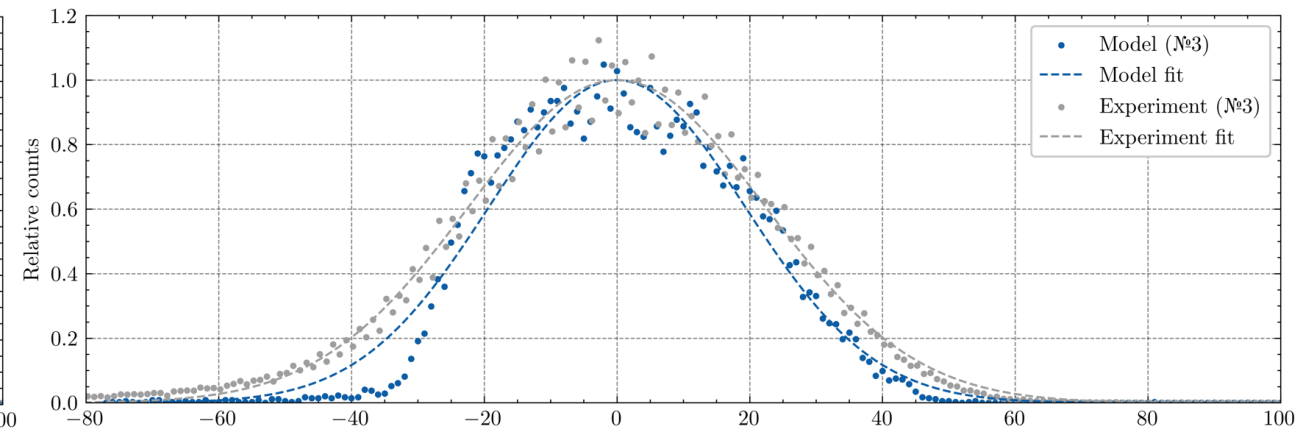
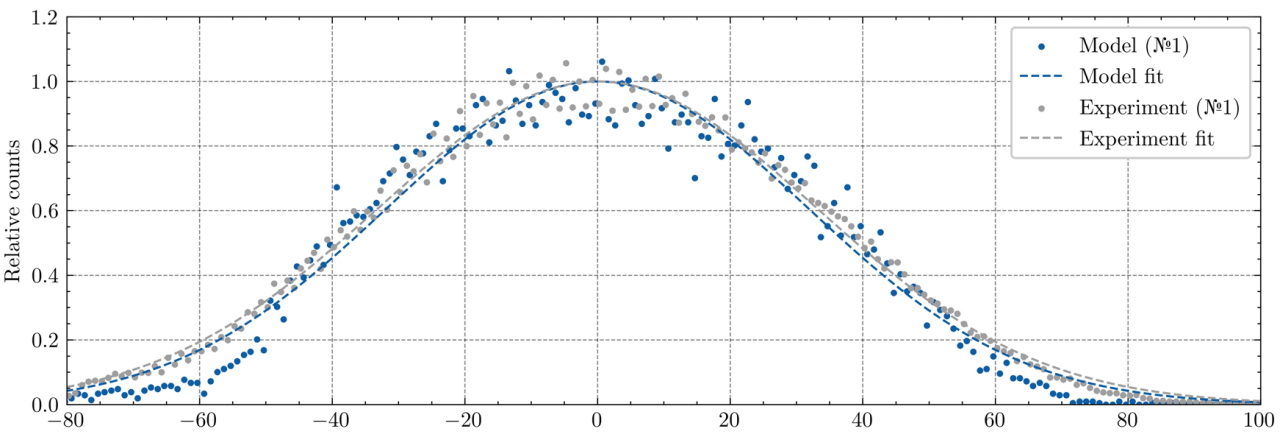
Comparison with the experiment



Comparison with the experiment



Time distributions of ODS triggers (0 equals central Carpet trigger time)



(x, y) and (ϑ, φ) reconstruction precision

